

LETTER TO THE EDITOR

The central region of spiral galaxies as seen by *Herschel**

M 81, M 99, and M 100

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ABSTRACT

With appropriate spatial resolution, images of spiral galaxies in thermal infrared ($\sim 10\ \mu\text{m}$ and beyond) often reveal a bright central component, distinct from the stellar bulge, superimposed on a disk with prominent spiral arms. ISO and *Spitzer* studies have shown that much of the scatter in the mid-infrared colors of spiral galaxies is related to changes in the relative importance of these two components, rather than to other modifications, such as the morphological type or star formation rate, that affect the properties of the galaxy as a whole. With the *Herschel* imaging capability from 70 to 500 μm , we revisit this two-component approach at longer wavelengths, to see if it still provides a working description of the brightness distribution of galaxies, and to determine its implications on the interpretation of global far-infrared properties of galaxies.

We quantify the luminosity of the central component by both a decomposition of the radial surface brightness profile and a direct extraction in 2D. We find the central component contribution is variable within the three galaxies in our sample, possibly connected more directly to the presence of a bar than to the morphological type. The central component's relative contribution is at its maximum in the mid-infrared range and drops around 160 μm to reach a constant value beyond 200 μm . The central component contains a greater fraction of hot dust than the disk component, and while the colors of the central components are scattered, colors of the disk components are more homogenous from one galaxy to the next.

Key words. infrared: galaxies – submillimeter: galaxies – galaxies: spiral – galaxies: nuclei – galaxies: fundamental parameters

1. Introduction

The central region of spiral galaxies hosts a number of processes that have little or no counterpart in the disks. A nuclear starburst, triggered by bar instabilities and associated or not with an active galactic nucleus (AGN) can often be found. The stellar density is also enhanced with respect to the disk, leading to an increase in the interstellar radiation field (ISRF) intensity. It is thus no surprise that in the mid-infrared (MIR), spiral galaxies can often be seen as the sum of a central component, hereafter referred to as the core component as distinct from the bulge, and a disk component (see e.g. Bendo et al. 2007). Extraction of the core component from MIR brightness profiles have revealed that it can contribute up to 90% of the total MIR luminosity and that the distribution of core/total luminosity ratio is evenly spread between 0 and 1 (Roussel et al. 2001a). This study also showed that much of the scatter in the MIR *global* properties of spiral galaxies came from changes in the relative contribution of the core component and of its spectral energy distribution (SED), while disk components formed a much more homogenous group. In

short, the scatter affecting the *global* MIR colors of galaxies is caused by a very *local* effect, i.e., the importance of the core component, rather than by a change in any single property of galaxies such as the ISRF throughout the disk. Taking this into account, Roussel et al. (2001b) shows that the MIR luminosity of galaxies could be turned into an accurate tracer of the star-formation rate, provided it was corrected for the core contribution. More recently, Chanial et al. (2007) extends this analysis by showing that the dispersion in the correlation between the infrared luminosity of galaxies and the dust temperature could be reduced by considering a compactness parameter, again related to the relative contribution of the core and disk components.

Care must therefore be taken when extrapolating locally established relations, such as the correlation of star-formation tracers in galactic disks, to global studies where galaxies are no longer resolved. One first needs to verify that the total luminosity of galaxies is indeed dominated by the component where these correlations have been established, or to design methods to correct for the existence of multiple contributors to the total luminosity. With its unprecedented sensitivity and spatial resolution, *Herschel* (Pilbratt et al. 2010) now allows us to carry this very investigation over the full thermal infrared peak, by combining the PACS instrument (Poglitsch et al. 2010), delivering

* *Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with significant participation from NASA.

Table 1. Relevant data for the galaxies studied here, as well as the images used (M: MIPS, P: PACS, S: SPIRE).

| Name | Type | Nuc ^a | D_{25} (') | Available data |
|-------|-------|------------------|--------------|---|
| M 81 | SAab | Sy 1.5 | 26.92 | M(24, 70, 160 μm), P(70, 160 μm), S(250, 350, 500 μm) |
| M 99 | SAC | HII | 5.37 | M(24, 70, 160 μm) S(250, 350, 500 μm) |
| M 100 | SABbc | T2 | 7.41 | M(24, 70, 160 μm) S(250, 350, 500 μm) |

Notes. ^(a) Nuclear classifications from [Ho et al. \(1997\)](#), T2 meaning a transition object between a starburst and a Sy 2.

spatial resolutions of 5''.5, 6''.7, and 11'' (*FWHM*) at 70, 100, and 160 μm , and the SPIRE instrument ([Griffin et al. 2010](#)) delivering spatial resolutions of 18''.1, 25''.2, and 36''.9 (*FWHM*) in its three photometric bands at 250, 350, and 500 μm .

2. The objects and data reduction methods

Galaxies studied here are part of the *Herschel* Reference Survey ([Boselli et al. 2010](#)), a SPIRE volume-limited imaging survey of 323 galaxies. Out of the 4 noninteracting spirals observed during the science demonstration phase, NGC 3683 and NGC 3982 are too small in apparent size for the present study, leaving M 99 and M 100 (see also [Pohlen et al. 2010](#)), to which we add M 81, which was observed as part of the Very Nearby Galaxies program ([Bendo et al. 2010b](#)). The properties of the objects that are relevant to this study are in Table 1.

Details of the SPIRE data reduction can be found in [Pohlen et al. \(2010\)](#) and [Bendo et al. \(2010b\)](#), while the PACS data reduction of M 81 is described in [Bendo et al. \(2010b\)](#). Photometric uncertainties are quoted at $\pm 15\%$ for all SPIRE bands and at $\pm 10/20\%$ for the PACS 70/160 μm bands by the instrument teams. We also use *Spitzer* MIPS maps at 24, 70, and 160 μm (see [Bendo et al. 2010a](#); [Clements et al. 2010](#), for data reduction details). The SPIRE maps are calibrated in Jy/Beam, and we used a beam size of 9.28×10^{-9} , 1.74×10^{-8} , and 3.57×10^{-8} sr at 250, 350, and 500 μm , respectively, to convert these into more conventional brightness units. The MIPS maps are calibrated in MJy/sr, and the PACS maps in Jy/pixel, with a pixel size of 1''.4 and 2''.85 for the 70 and 160 μm bands.

3. Extraction of the core components

Our first approach uses radial surface brightness profiles measured inside concentric ellipses. For all maps of a given galaxy, we used the same parameters to build the integrating ellipses, namely the central coordinate, the ellipse axis ratio, and the position angle; i.e., these parameters are not fit to the data but obtained from NED¹. This is a coarser approach than the one used by [Pohlen et al. \(2010\)](#), where the ellipse parameters are first fitted to the maps before the profile integration is made. A simpler approach is justified here since we are not searching for the most accurate structural description of the surface brightness distribution in galaxies, but rather for a method to systematically measure the contribution of the central region to the total galaxy luminosity. Errors in the computed average surface brightness are the quadratic sum of the error on the mean brightness, of the error derived by accounting for the corresponding

error map on the same ellipses, and of the error on the background level. Generally, the total errors are much smaller than the computed brightness except close to the background. For all maps, we intentionally went significantly beyond the D_{25} diameter. To obtain the galaxy's brightness profile, we determined the background brightness as the mean of the brightness measured typically about 100'' beyond the point where the galaxy emission ceased to be detected.

To search for and extract the core component, we obtained the best (in the least-square sense) combination of Gaussian or exponential functions that represent the measured radial profiles. Expressed in units of surface brightness as a function of radius, we fit the following function to the data:

$$S(r) = \sum_{i=1}^N S_i^0 \exp - \left(\frac{r}{r_i^s} \right)^{k_i}, \quad (1)$$

with S^0 the central surface brightness, r^s the scale length, $k = 1, 2$ for exponential or Gaussian functions², and $N = 1$ or 2. The fitting function was smoothed by the appropriate PSF, so the parameters we derived describe the intrinsic light distribution. It is important to recall that our aim is to build a robust and systematic method to get at the one parameter we are searching for: the relative contribution of the core component to the total flux, and not to achieve the best representation of the full profile.

The use of either Gaussian or exponential functions is a result of the fit, but that some of the disk profiles prefer a Gaussian to an exponential must not be overinterpreted because the presence of prominent spiral arms can create a distinctive bump in the disk part of the profile that is represented by a Gaussian better than by an exponential. In either case we did not use the fit parameters that describe the "disk" component in the rest of the analysis. The central component is almost always represented by a Gaussian profile very well, with the disk component also adequately fit at the location where the core profile merges into the disk profile (see Fig. 1). We then decomposed the luminosity of the galaxies in the following way. The total luminosity was obtained by integrating the measured radial profile up to the background radius, the total luminosity of the core component was obtained by integrating the fitted profile of the core, and the disk luminosity is simply the difference between these two quantities. In Table 2 we list the core components' scalelength and contribution to the total luminosity derived from the fits.

In a second approach, we performed 2D fitting on M 81 (chosen because of the level of details available and the existence of a point-like component in the center), using the GALFIT code ([Peng et al. 2002](#)). Galactic structure (arms, bars, rings) is much more complex to represent mathematically in 2D, thus this method converges on the extraction of unresolved components. This provides a cross-check to the radial profile approach since, when the central component is resolved in the profile, the 2D method should underestimate it, but both methods should agree when it is unresolved. This is what we observe (see below).

4. Discussion

In general, we observe that the same family of functions provides the best-fit models to a given galaxy, regardless of the wavelength. This is mostly because the same morphological features dominate the profile of galaxies over the explored spectral domain, a fact worth noting given the quite drastic changes in the

¹ <http://nedwww.ipac.caltech.edu/>.

² In the case of a Gaussian, we have $r^s = 0.6 \times FWHM$.

Table 2. Scale-lengths and relative contributions of the core component.

| Band (1) | M 81 | | | M 100 | | |
|---------------------|------------------|------------|--------------|------------------|------------|--------------|
| | r^s (") (2) | Fit (3) | f_c (4) | r^s (") (5) | Fit (6) | f_c (7) |
| M 24 μm | 7.2 | E+E | 8.8 | 10.3 | G+G | 20.1 |
| M 70 μm | 38.8 | G+G | 6.2 | 13.2 | G+G | 12.2 |
| P 70 μm | 27.9 | G+G | 6.3 | <i>n.a.</i> | | |
| M 160 μm | 28.5 | G+G | 0.3 | 5.14 | G+G | 3.5 |
| P 160 μm | 26.4 | G+G | 1.2 | <i>n.a.</i> | | |
| S 250 μm | 25.2 | G+G | 0.3 | 9.5 | G+G | 7.8 |
| S 350 μm | 10.2 | G+G | 0.2 | 5.4 | G+G | 7.4 |
| S 500 μm | 11.8 | G+G | 0.2 | 6.5 | G+G | 7.6 |

Notes. For M 99, the profile fit revealed no central component except at 24 μm so we report this in the text. In column (1), M, P and S are MIPS, PACS and SPIRE, respectively, (2) and (5) the core component scale-length in arcseconds, (3) and (6) the best fit functions, E for exponential and G for Gaussian, and (4) and (7) the core component contribution to the total flux in %, respectively for M 81 and M 100. Relative uncertainties on the scale parameters are typically of 10%, while on the core fractions these are typically 20%

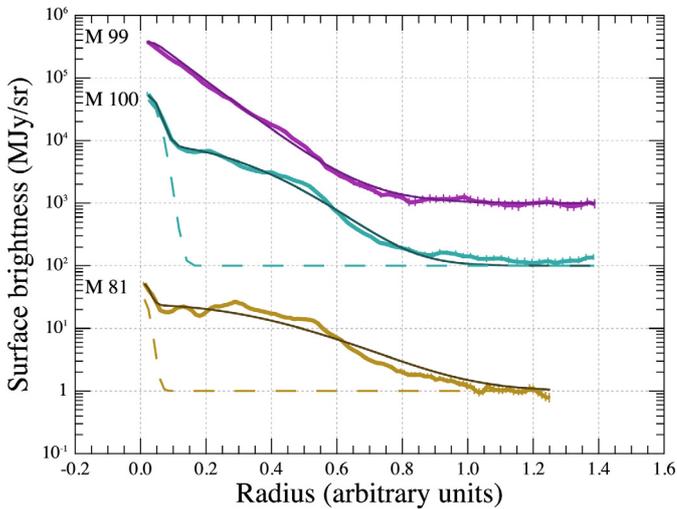


Fig. 1. The surface brightness radial profiles at 250 μm (thick curves with error bars) and the fitted profiles. We also plot as dashed lines the core component fit, to visualize its relative contribution. A value of 1 was added to all profiles to avoid visually enhancing the noise in the outer part of the profile, and the M 99 and M 100 profiles are multiplied by 10^2 and 10^3 for clarity. The abscissa of the graph is the radius normalized by the distance at which we measure the background.

thermodynamical state and physical nature of the mid to far IR-emitting grains, from out-of-equilibrium, nearly molecular-sized grains to thermalized sub-micron grains (but see Pohlen et al. 2010, for a discussion of the full profiles properties). However, the profiles already show that the relative importance of the core component is variable from one galaxy to the next (see Fig. 1).

M 81 - This is the closest galaxy and thus its central region is observable with a high level of detail (see Bendo et al. 2010b). The core is compact, possibly the IR counterpart of the Seyfert nucleus, on which a bright arc, presumably the start of one of the spiral arms, connects. Moving away from the center, the brightness of the arm diminishes until it increases again to form a bright, almost ring-like structure. Thus all profiles, except the 24 μm one, show a first local minimum at $\sim 80''$ radius (1.4 kpc at 3.65 Mpc). With high enough resolution, the double nature of the core component (central unresolved source and spiral arm) is visible in the profile as a break in the slope of the core profile.

The interplay between the resolved and unresolved parts of the core is clear in the comparison of the GALFIT and profile-fitting core-fraction estimates: as the former preferentially extracts the unresolved component, while the latter is able to extract both the unresolved and resolved components as a single structure departing from the disk, we expect both methods to agree, for a given band, when either the unresolved source is dominant or the whole core structure is unresolved. This is indeed observed for MIPS 24 μm (where the unresolved component dominates) and for MIPS 160, SPIRE 350, and 500 μm (where the whole core structure is not, or barely, resolved). In the other bands, the unresolved component is less dominant, and GALFIT computes core fractions that are a factor of 3 below those coming from the profile fitting. However, inspection of the corresponding images clearly shows the unresolved component sitting on a core structure distinct in brightness from the disk. The core contribution derived with the profile fitting method is also very compatible with what can be measured naively from the profile itself, inside the $80''$ radius. The core component is generally well represented by Gaussians and, when resolved, the derived scalelengths are compatible with one another, given typical uncertainties of 10%. We note a severe disagreement between the two 160 μm core fraction measurements, but inspection of the profiles clearly shows that this stems from dilution of the barely resolved central component in the disk emission at the low resolution of *Spitzer*, thereby pointing to the intrinsic limitations of the MIPS 160 μm band as the way to perform these decomposition studies. At increasing SPIRE wavelengths, we observe as well that the central component becomes gradually unresolved.

M 99 - Although this galaxy clearly shows a bright center and spiral arms in all available IR bands (see Pohlen et al. 2010), our fit of the profiles is usually adequate with a single exponential function; however: (1) an exponential does not accurately represent the disk profile that has a distinct break (at $r \simeq 0.4$ in Fig. 1, see Pohlen et al. 2010, for a discussion of this feature), and (2) the 24 μm image is sharp enough to reveal a barely resolved Gaussian component in the center ($r^s = 3''$) contributing less than 0.5% of the total 24 μm flux.

M 100 - The barred nature of the galaxy clearly shows up as a bump in the high-resolution profiles (e.g. 24 μm and 250 μm , see also Pohlen et al. 2010), and as a result, the fitting function favors a double Gaussian profile, with very good fits of the central component achieved from 24 μm to 500 μm . Table 2, however, shows the central component is unresolved at 160, 350, and 500 μm , while for the other wavelengths a scale-length of $11'' \pm 2''$ is indicated by the fits. When the central component is unresolved, our capacity to accurately measure its contribution will depend on how bright it is compared to the disk component. Thus it is possible that the core fraction we compute at 160 μm (the band with the worst resolution) is in fact a lower limit to the actual core fraction, pointing again to the limited ability of the MIPS 160 μm band to accurately measure core contributions.

The necessity to account for a core component to represent the surface brightness profile already appears weakly related to the morphological type: it is on the intermediate type galaxy M 100 that we observe the largest core fraction. It is quite tempting, and probably logical, to relate that to the presence of a weak bar in M 100, but this can only be asserted on a reasonable sample, as the presence of a bar does not guarantee there is a strong MIR core component (Roussel et al. 2001a).

Interestingly, even though M 81 and M 100 have significantly different core fractions (by a factors >2), we observe in Fig. 2 the same trend in the evolution of this fraction with wavelength: it

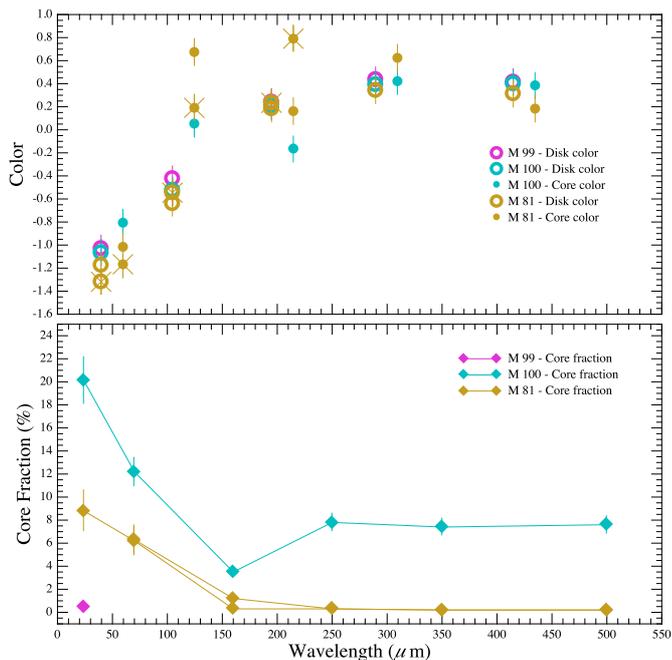


Fig. 2. *Top panel:* colors as a function of wavelength. Colors are displayed at an abscissa in-between the band wavelengths used, e.g., the $[24]/[70]$ color is plotted at $50\ \mu\text{m}$, with open symbols for the disk colors and filled symbols for the core colors. For M81, we identify the colors incorporating PACS data with an extra cross overlaid on the corresponding symbol. For clarity, the core and disk colors are offset in x by $10\ \mu\text{m}$. *Bottom panel:* Core fractions as a function of wavelength, again with two series of data for M81, with MIPS and PACS. Error bars can be smaller than symbol size.

is maximum at the shorter wavelengths and nearly constant beyond $200\ \mu\text{m}$, with the transition occurring in the $70\text{--}200\ \mu\text{m}$ range. The evolution of the core fraction can be interpreted by inspecting the component's colors. We show in Fig. 2 the colors of the core and disk components built from neighboring bands. This figure first shows quite clearly that all the disk colors are very similar for the three galaxies studied here, even when the core component represents a substantial fraction of the emission ($10\text{--}20\%$ at $\lambda \leq 160\ \mu\text{m}$). It also shows that only the $[70]/[160]$ ³ core colors, and to a lesser extent $[160]/[250]$, significantly differ from the disk colors. The $[70]/[160]$ core colors are systematically higher than those of disks, showing the core emission as coming from hotter dust than the disk emission. If the dust is hotter in the core than in the disk, then it will represent a larger fraction of the emission at shorter wavelengths. In the SPIRE range, we reach the Rayleigh-Jeans regime, where colors have no dependence on temperature. That color differences are more prominent in the $70\text{--}250\ \mu\text{m}$ range is simply because this is where the peak of the thermal emission of dust in galaxies occur, showing that this range is key to properly understanding the infrared emission of galaxies.

To reconcile the small $[24]/[70]$ color variation between disks and cores with the larger core fractions at $24\ \mu\text{m}$, we suggest that it is because the $24\ \mu\text{m}$ and some of the $70\ \mu\text{m}$ emission come from transiently heated grains, i.e., show little spectral variation with the ISRF intensity or hardness. A harder ISRF in the core could then explain the larger core fractions observed at $24\ \mu\text{m}$ by boosting the emission from transiently heated

grains without significantly affecting their emission spectrum. The small AGN in M81 is a natural explanation for this, as is a possible bar-driven starburst in the center of M100. We also tested an alternative explanation in M81, namely a direct contribution of bulge stars in the $24\ \mu\text{m}$ band, as suggested by Bendo et al. (2008), by correcting the MIPS image using IRAC data (a correction to the total flux of -5%). The “dust-only” core fraction is almost unchanged, revealing that bulge stars are not responsible for the FIR core structure.

5. Conclusions

In this small set of galaxies, the core fractions are small ($\leq 20\%$), therefore the correction to the full SED of the galaxies introduced by accounting for the different core properties would be minimum. Nevertheless we already observe both a significant variation in the core component contribution from one galaxy to the next and a systematic decrease with wavelength. We also find that the galactic disks studied here show remarkable color homogeneity over the $24\text{--}500\ \mu\text{m}$ range, while the core components are notably hotter as shown by their higher $[70]/[160]$ color value. This opens the possibilities that (i) across the *Herschel* wavelength range, the global SED of a galaxy may be influenced as much by the relative importance of the core and disk components as by a global property, such as the star formation rate or the morphological type, and (ii) that this may depend on the observed wavelength. We shall thus examine this in larger samples such as the HRS set of resolved spiral galaxies.

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³ Shorthand for $\log [f_\nu(70\ \mu\text{m})/f_\nu(160\ \mu\text{m})]$.

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